ATOMS OF SPACETIME

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Fourth Abdus Salaam Memorial Lecture (2005-2006)

WHAT WILL BE THE VIEW REGARDING GRAVITY AND SPACETIME IN THE YEAR 2206 ?

CLASSICAL GRAVITY = CONDENSED MATTER PHYSICS OF SPACETIME SUBSTRUCTURE

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HOW GRAVITY RESPONDS TO AND MODIFIES THE STRUCTURE OF THE VACUUM IS PROBABLY THE KEY QUESTION IN THEORETICAL PHYSICS TODAY. A WORD FROM THE SPONSORS

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"I do not think we have a completely satisfactory relativistic quantum mechanical theory; even one, that does not agree with nature but at least agrees with logic. Therefore, I think the renormalisation theory is simply a way to sweep the difficulties of the divergences under the rug."

R.P.Feynman [1965, Nobel Prize lecture]

Non Relativistic Quantum Mechanics



In NRQM a path is [X(T), Y(T), Z(T)]. Paths go backward in space X, Y, Z but not in time T

Relativistic Quantum Mechanics



In RQM a path is [T(s), X(s), Y(s), Z(s)] where s is the proper time. Paths go backward in space X, Y, Z and time T !



Space



Space



Space



Space



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VACUUM DEPENDS ON THE SCALE OF PROBING

• Phonons, magnons, photons, electrons, protons, quarks Each can arise as an excited state of a suitably defined system of oscillators.



• Example: Lattice vibrations when quantized will give rise to phonons. The ground state is that of zero phonons.

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• Ground state of electromagnetic field:

$$P \propto \exp\left[-rac{1}{16\pi^3\hbar c}\intrac{\mathbf{E}(\mathbf{x})\cdot\mathbf{E}(\mathbf{y})}{|\mathbf{x}-\mathbf{y}|^2}d^3xd^3y
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Key Point: You can never make the fluctuations go away.

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- Two conducting plates, kept in vacuum separated by *a* attracts each other with a force

$$\frac{F}{A} = -\frac{\pi^2 \hbar c}{240 a^4}.$$

The energy of the configuration is

$$\frac{E}{A} = -\frac{\pi^2 \hbar c}{720 a^3}; \qquad F = -\frac{dE}{da}$$

VACUUM STATE HAS RANDOM FIELD MODES



SOME OF THESE MODES VIOLATE BOUNDARY CONDITIONS AT PLATES



NEW VACUUM STATE IS DIFFERENT FROM THE ORIGINAL ONE!



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With plates : $k = (2\pi/a)n, n = 0, 1, 2, ...$ Without plates : $k = (2\pi/a)n, 0 < n < \infty$

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$$\Delta E = \frac{1}{2}\hbar \frac{2\pi c}{a} \left[\sum_{n} n - \int_{0}^{\infty} n dn \right] = -\frac{\pi \hbar c}{24a}$$

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• Pattern of vacuum fluctuations change when external conditions change.

DOES VACUUM GRAVITATE ?








IS VACUUM DOMINATING OUR UNIVERSE ?

- Our Universe has a totally preposterous composition which no cosmologist wanted !
- $\Omega_{\rm total} \approx 1$
 - $\Omega_{
 m radiation} \approx 0.00005$ $\Omega_{
 m neutrinos} \approx 0.005$ $\Omega_{
 m baryons} \approx 0.04$ $\Omega_{
 m wimp} \approx 0.31$
 - $\Omega_{\rm darkenergy} \approx 0.65$
- What is this Dark Energy ?





Newtonian physics: A can communicate with B and C.



Special Relativity introduces a causal horizon; A cannot communicate with B.



A family of observers $(O_1, O_2, O_3, ...)$ have a causal horizon



$$ds^{2} = -\left(1 - \frac{2M}{r_{s}}\right) dt_{s}^{2} + \left(1 - \frac{2M}{r_{s}}\right)^{-1} dr_{s}^{2}$$

= $Q_{s}^{2}(T, X)(-dT^{2} + dX^{2})$



$$ds^{2} = -(1 - H^{2}r_{dS}^{2}) dt_{dS}^{2} + (1 - Hr_{dS}^{2})^{-1} dr_{dS}^{2}$$

= $Q_{dS}^{2}(T, X)(-dT^{2} + dX^{2})$



 $ds^2 = -2\kappa x_{\scriptscriptstyle R} dt^2 + (2\kappa x_{\scriptscriptstyle R})^{-1} dx_{\scriptscriptstyle R}^2$ = $-dT^2 + dX^2$

KEY NEW FEATURE IN COMBINING GRAVITY AND QUANTUM THEORY

VACUUM STATE IS OBSERVER DEPENDENT!





 $\phi(T, X) = \exp[-i\Omega(T - X)]$



$$\phi(T,X) \equiv \phi\left(T(\tau), X(\tau)\right) = \exp{-i\Omega\left(rac{1-v}{1+v}
ight)^{1/2} au}$$

Doppler effect: $\Omega' = \Omega\left(rac{1-v}{1+v}
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• A mode $\phi(t,x) = \exp[-i\Omega(t-x)]$ frequency Ω will lead to

$$\phi(\tau) = \exp[i\Omega a^{-1}\exp(-a\tau)] = \int_0^\infty d\nu [A(\nu)e^{-i\nu t} + B(\nu)e^{i\nu t}]$$

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- For a black hole, $a = GM/R^2$ at $R = 2GM/c^2$. This gives $k_BT = \hbar c^3/8\pi GM$, the Hawking temperature.

• Metric:

$$ds^{2} = f(r)dt^{2} - f(r)^{-1}dr^{2} - r^{2}(d\theta^{2} + \sin^{2}\theta d\phi^{2}) \equiv f(r)dt^{2} - f(r)^{-1}dr^{2} - dL_{\perp}^{2}$$

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$$(L_P^2 \equiv G\hbar/c^3)$$
:

$$S = \frac{1}{4L_P^2} (4\pi a^2) = \frac{1}{4} \frac{A_H}{L_P^2}; \quad E = \frac{c^4}{2G} a = \frac{c^4}{G} \left(\frac{A_H}{16\pi}\right)^{1/2}$$

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• Multiply *da* to write:



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• In normal units, there is still no \hbar in TdS!

HOLOGRAPHIC DUALITY OF EINSTEIN GRAVITY T.P.,gr-qc/0311036; gr-qc/0412068

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allowing a dual description of gravity using either L_{bulk} or L_{sur} !
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• We can obtain dynamics of gravity using *only* the surface term of the Hilbert action.

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TdS - pdV - dE = 0

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SAKHAROV PARADIGM

Gravity as an Emergent Phenomenon

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<u>SOLIDS</u>

SPACETIME

Mechanics; Elasticity $(\rho, \mathbf{v} \dots)$

Einstein's Theory $(g_{ab} \dots)$

Statistical Mechanics

of atoms/molecules

Statistical mechanics

of "atoms of spacetime"

SAKHAROV PARADIGM

Gravity as an Emergent Phenomenon

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SAKHAROV PARADIGM Gravity as an Emergent Phenomenon

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Einstein's Theory $(g_{ab} \dots)$

Thermodynamics of spacetime

Statistical mechanics

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KEY NEW RESULT: AREA QUANTISATION

In semi-classical limit, demanding $\exp iA_{sur} = \exp 2\pi in$ leads to area quantization:

Area =
$$(8\pi L_P^2)n$$

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$$\rho_{\rm vac} = \frac{\hbar c}{L_P^4} + \frac{\hbar c}{L_P^4} \left(\frac{L_P}{L_H}\right)^2 + \frac{\hbar c}{L_P^4} \left(\frac{L_P}{L_H}\right)^4 + \cdots$$

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$$\rho_{\rm de}\approx \sqrt{\rho_{\rm uv}\rho_{\rm ir}}\approx \frac{\hbar c}{L_P^2 L_H^2}\approx \frac{c^2 H^2}{G}$$

$$\rho_{\rm vac} = \frac{\hbar c}{L_P^4} + \frac{\hbar c}{L_P^4} \left(\frac{L_P}{L_H}\right)^2 + \frac{\hbar c}{L_P^4} \left(\frac{L_P}{L_H}\right)^4 + \cdots$$
'vacuum renormalisation'
makes this zero (?)

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REFERENCES

- T. Padmanabhan, The Holography of gravity encoded in a relation between Entropy, Horizon Area and the Action for gravity, Gen.Rel.Grav., 34 2029-2035 (2002) [gr-qc/0205090] [Second Prize essay; Gravity Research Foundation Essay Contest, 2002]
- T. Padmanabhan, Classical and Quantum Thermodynamics of horizons in spherically symmetric spacetimes, Class.Quan.Grav. 19, 5387 (2002).
 [gr-qc/0204019]
- T. Padmanabhan, Why gravity has no choice: Bulk spacetime dynamics is dictated by information entanglement across horizons, Gen.Rel.Grav., 35, 2097-2103 (2003) [Fifth Prize essay; Gravity Research Foundation Essay Contest, 2003]
- T. Padmanabhan, Gravity and the Thermodynamics of Horizons, Phys. Reports, 406, 49 (2005) [gr-qc/0311036]
- T. Padmanabhan, Holographic Gravity and the Surface term in the Einstein-Hilbert Action, Plenary talk at DICE-2004, Brazilian Jour.Phys. (Special Issue) 35, 362 (2005) [gr-qc/0412068]